

Linear and Circular Polarization in Diffuse Interstellar Bands

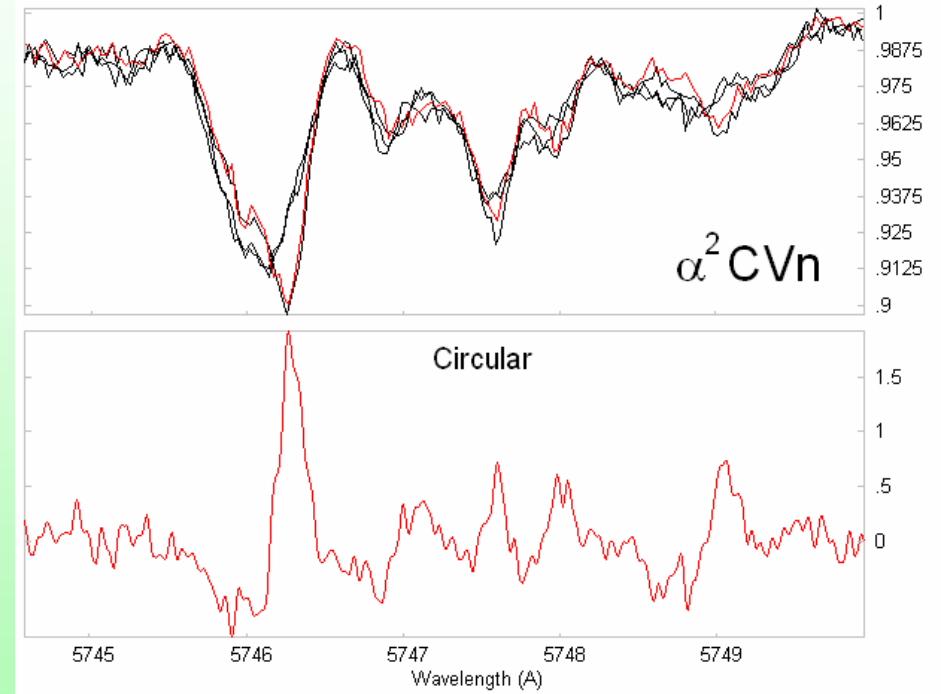
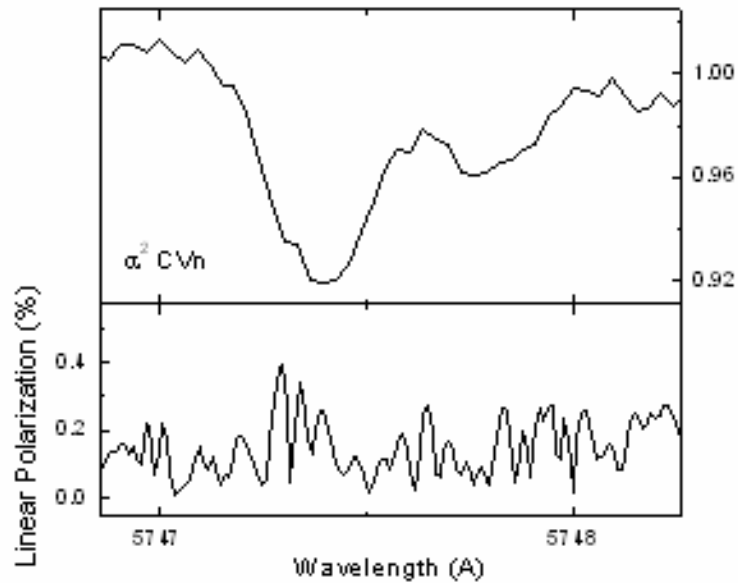
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KASI, Daejeon

Bohyunsan Optical Astronomy Observatory

- 1.8-m telescope
- Fiber feed Echelle Spectrograph
3,500 - 10,000 Å, $R = 35,000 - 90,000$



Polarization



Polarimetry of DIBs (1972 – 2007)

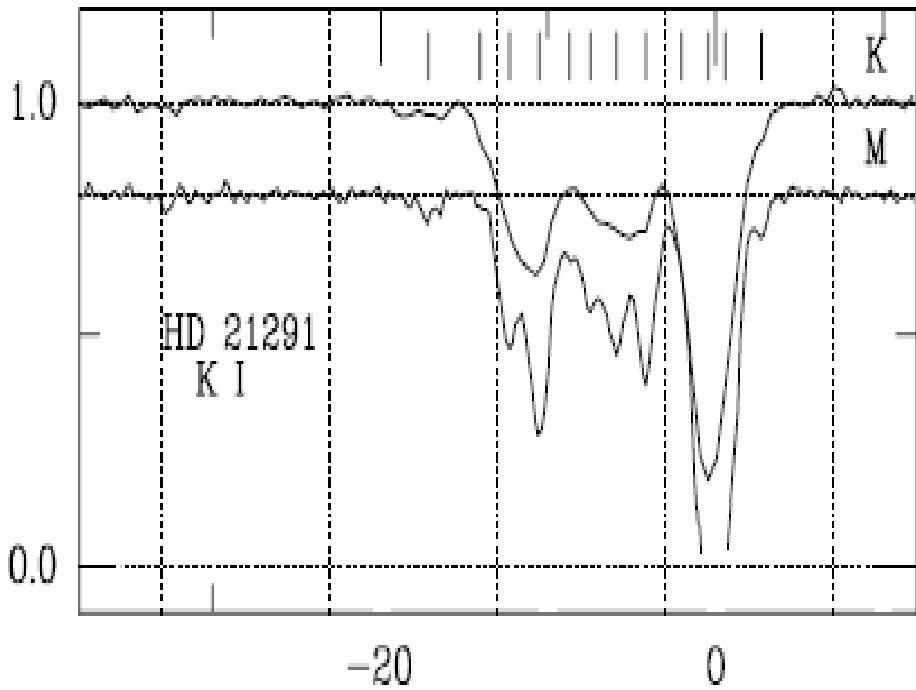
	Sp	E(B-V)	<i>DIBs</i>		
• HD 13854	B1Iabe	0.5			
• HD 21219	B9Ia	0.4			
• HD 21389	A0Iab	0.55			
• HD163472	B3V	0.25	<i>4430</i>	<i>5780</i>	<i>5797</i>
• HD183143	B7Ia	1.25	<i>6196</i>	<i>6203</i>	<i>6284</i>
• HD194279	B1.5I	1.2	<i>6376</i>	<i>6379</i>	<i>6613</i>
• HD197770	B2IV	0.55			
• HD198478	B3Ia	0.55			

BOAO data

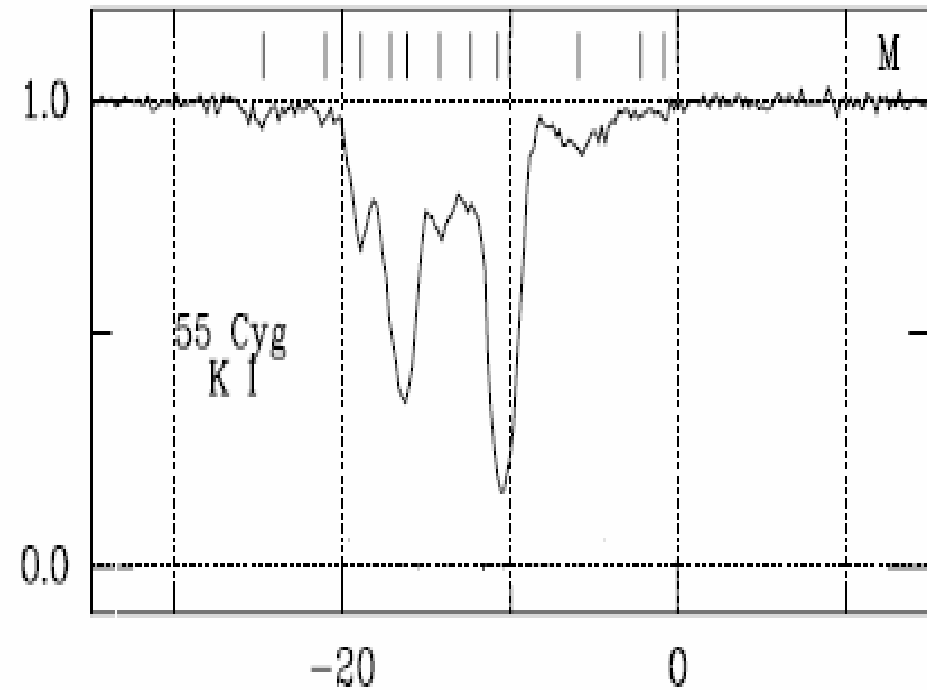
	Sp	E(B-V)
• HD 23180	B1III	0.28
• HD 24398	B1Iab	0.32
• HD 34078	O9.5V	0.5
• HD143275	B0.2IV	0.33
• HD144217	B0.5V	0.31
• HD147889:	B2III/IV	1.06
• HD226868:	O9.7Iab	1.03

+16 stars

HD 21291



HD198478



Welty & Hobbs, 2001, ApJS, 133, 345

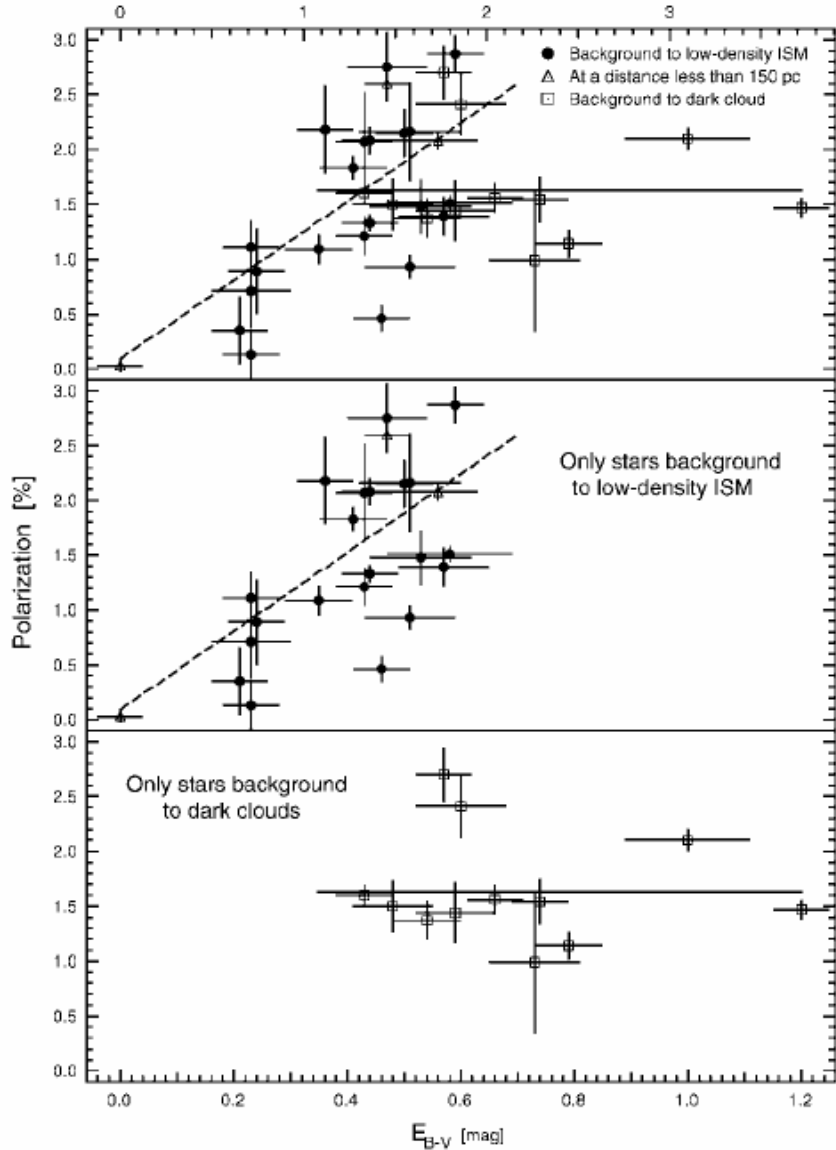


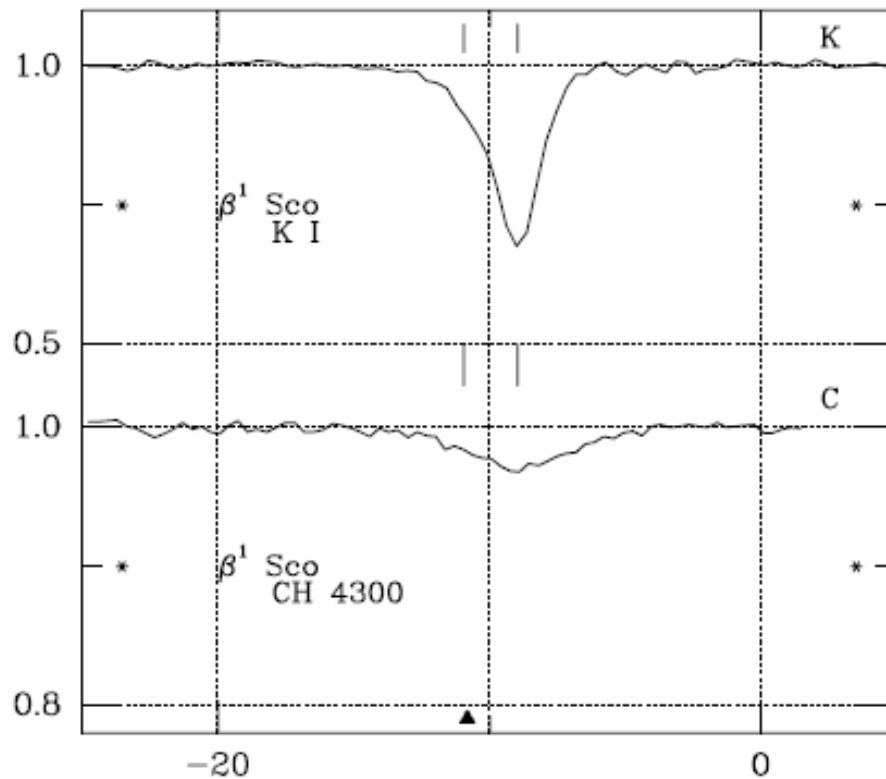
FIG. 2.—Observed relation between polarization and extinction, where $A_V = 3.1E_{B-V}$. The lines are least-square linear fits, weighted by the uncertainty in p . The dashed line is the fit to the points representing the stars background to the low-density ISM, which gives $p = (0.09 \pm 0.06) + (3.58 \pm 0.13)E_{B-V}$, with a correlation coefficient of 0.68. The solid line is the fit to the points representing the stars background to dark clouds, which gives $p = (1.61 \pm 0.13) + (0.03 \pm 0.15)E_{B-V}$, with a correlation coefficient of 0.79.

The polarizing power of dust

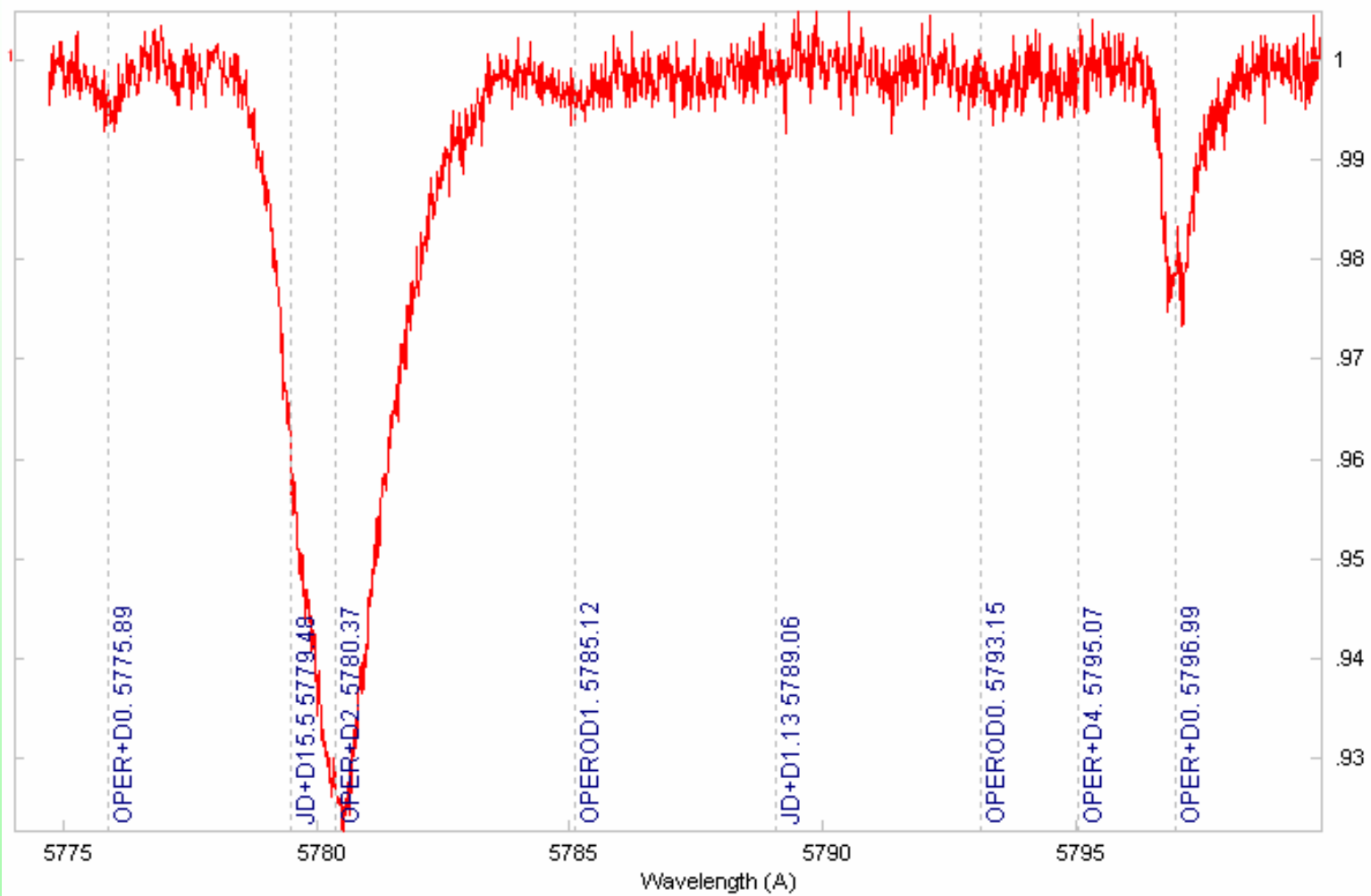
Arce et al. , 1998, ApJ, 499, L93

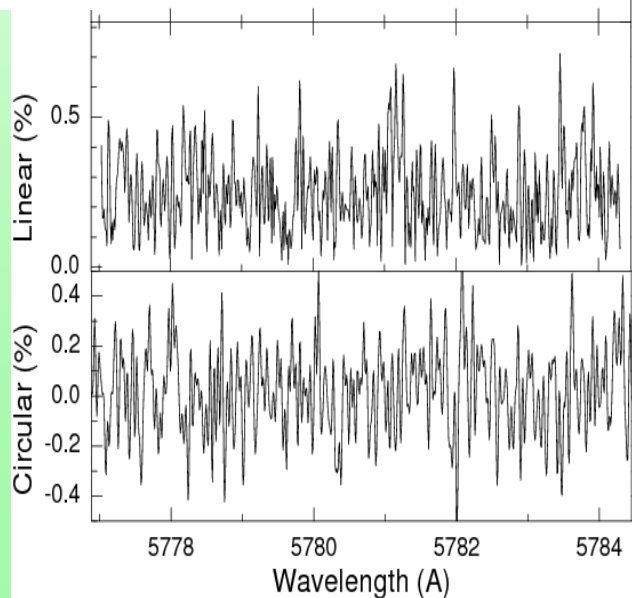
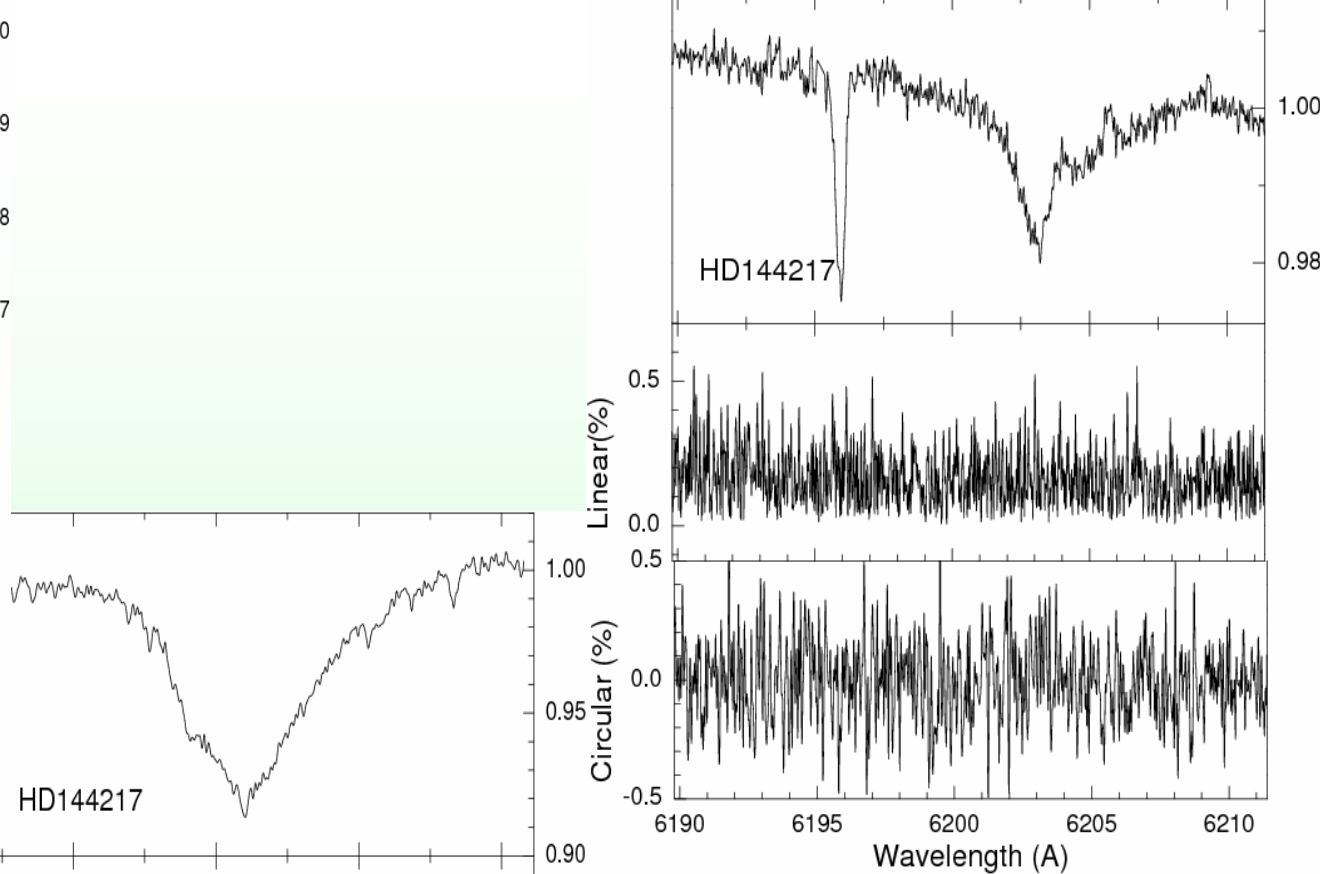
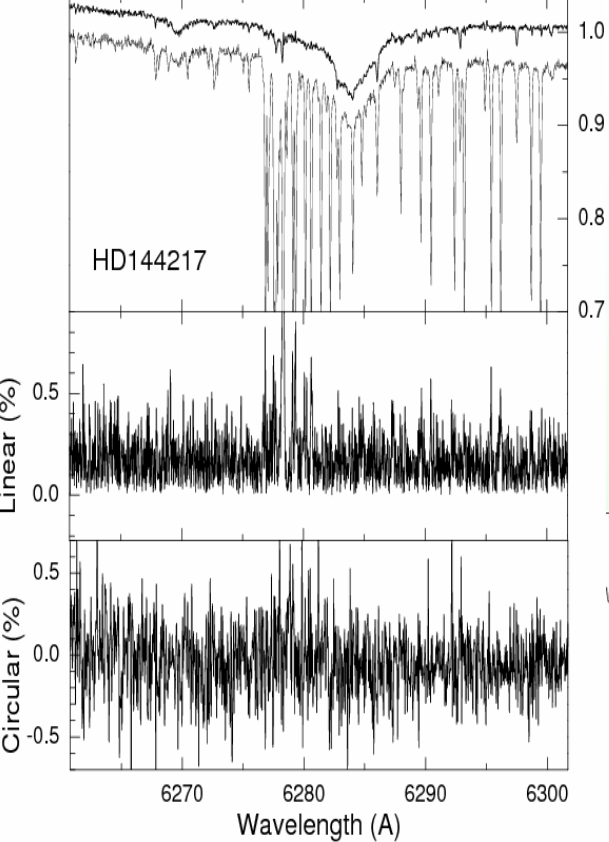
- 36 stars - Taurus dark clouds

HD144217

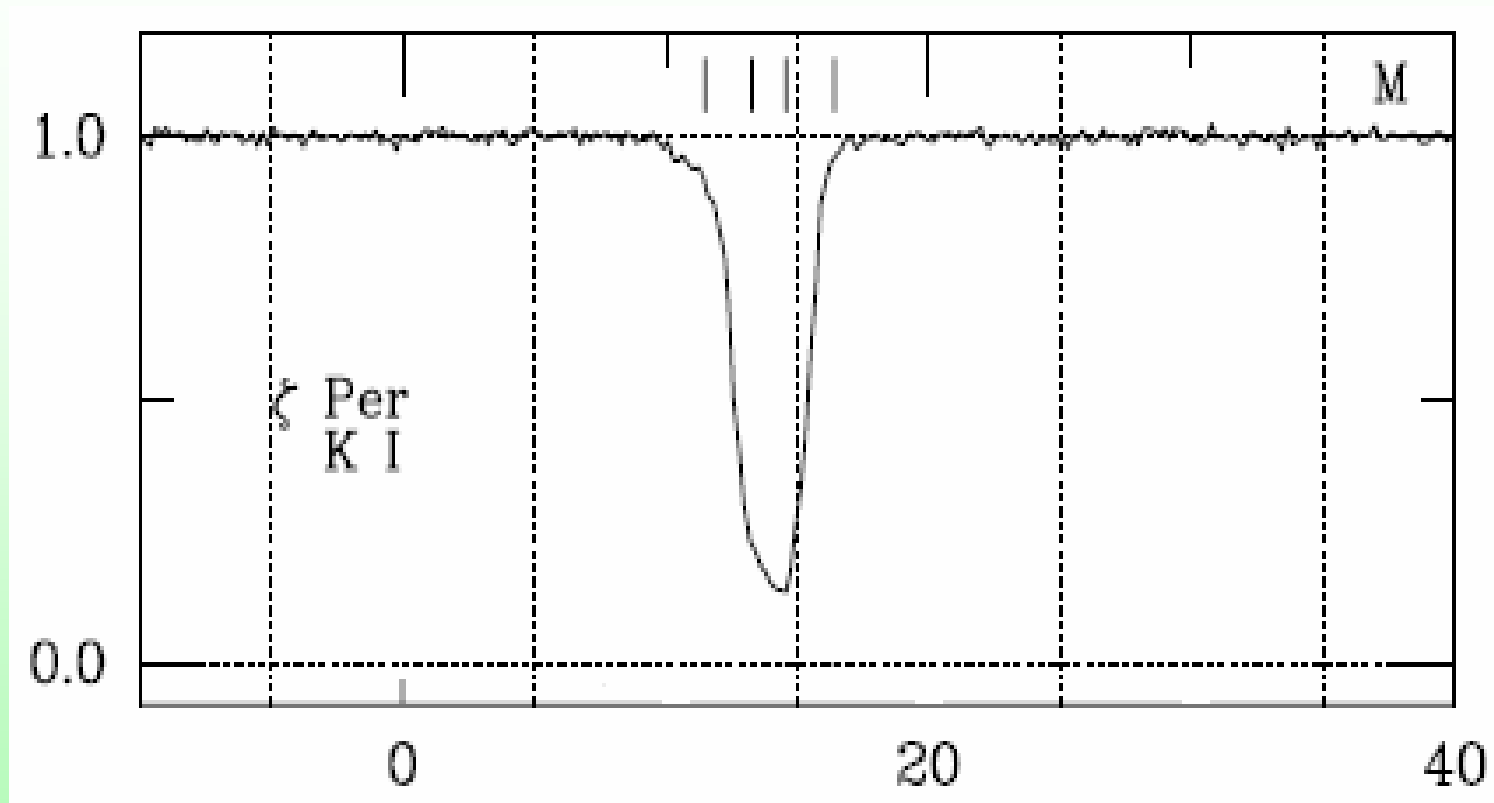


Welty & Hobbs, 2001, ApJS,133,345
R~200,000



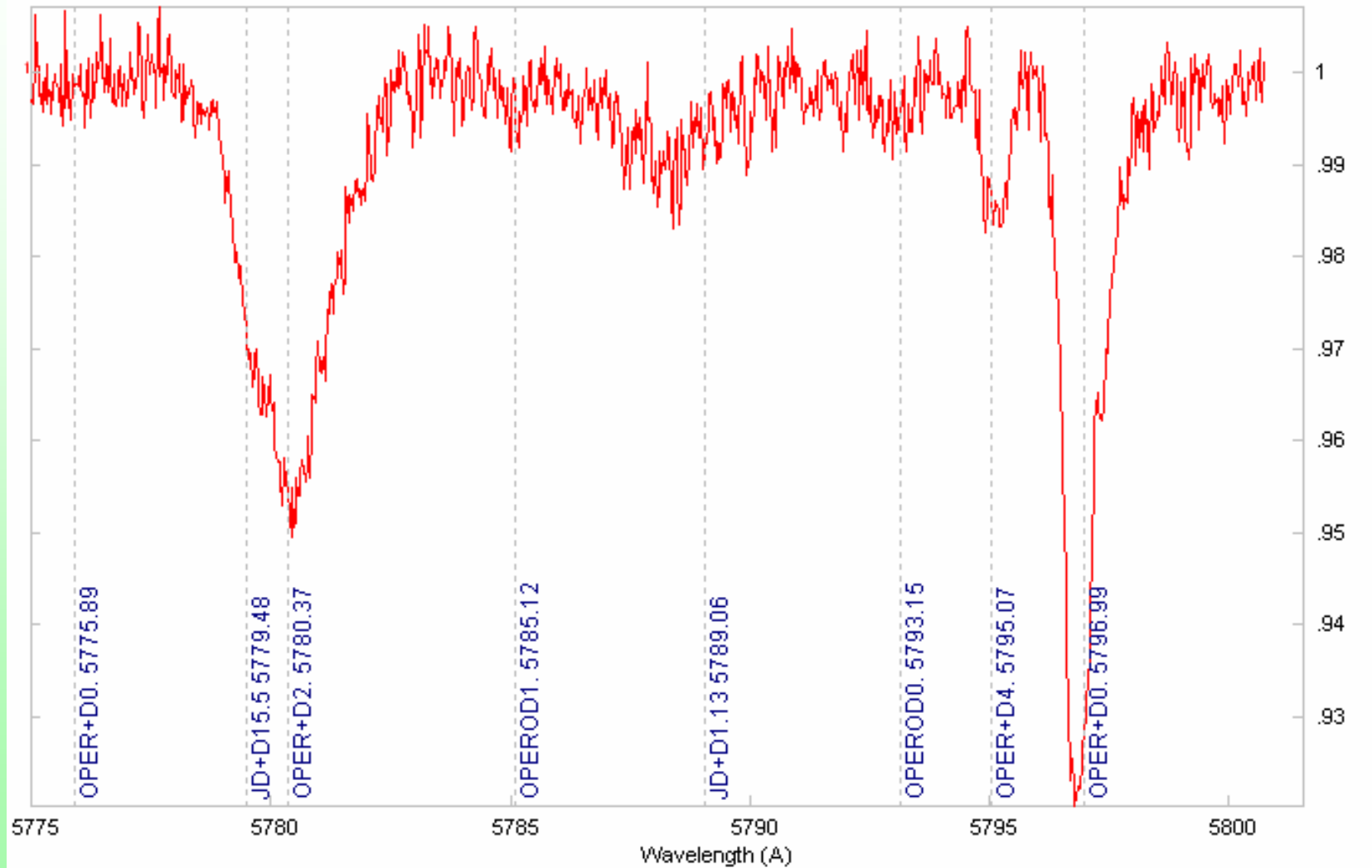


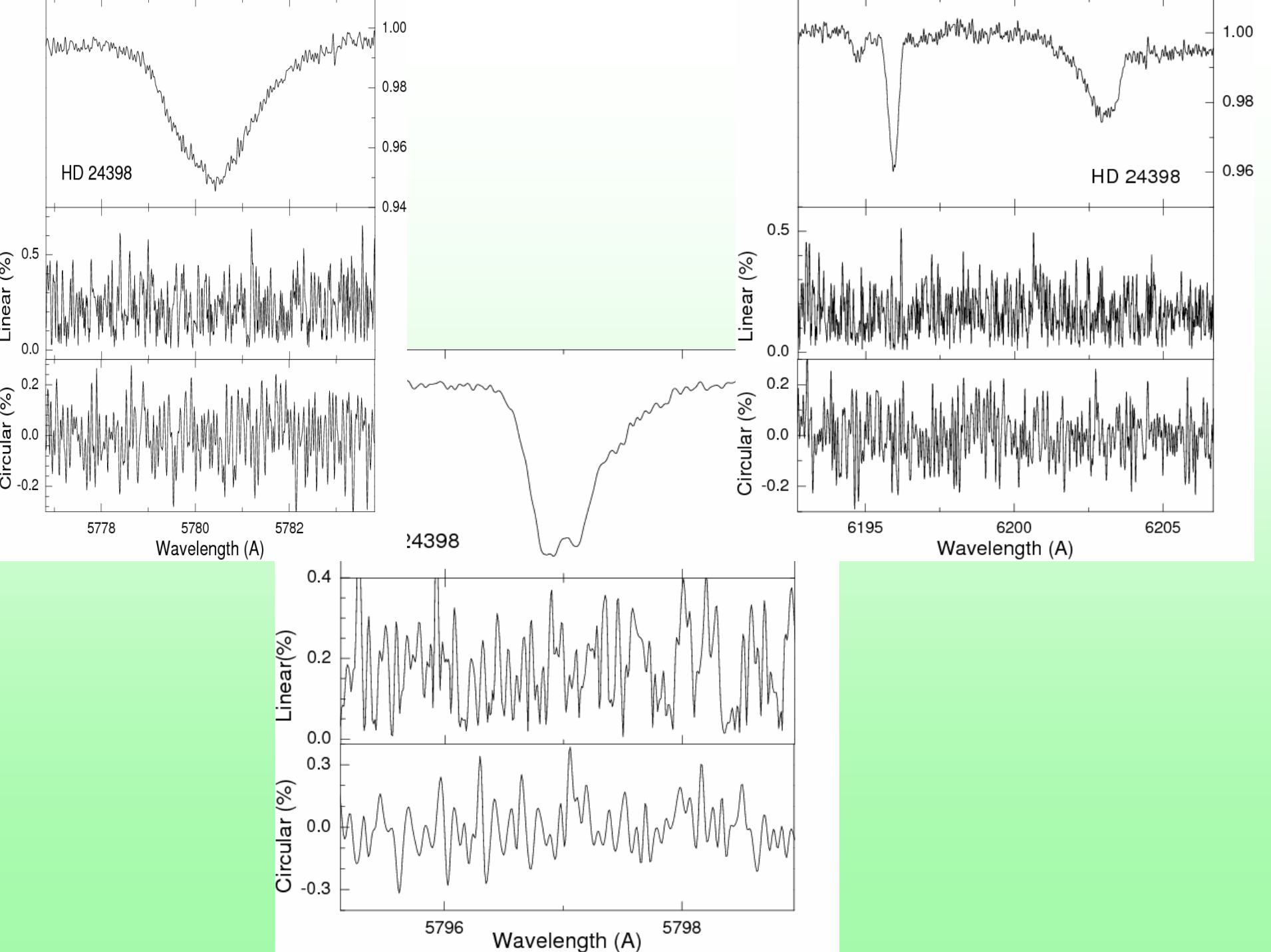
HD 24398

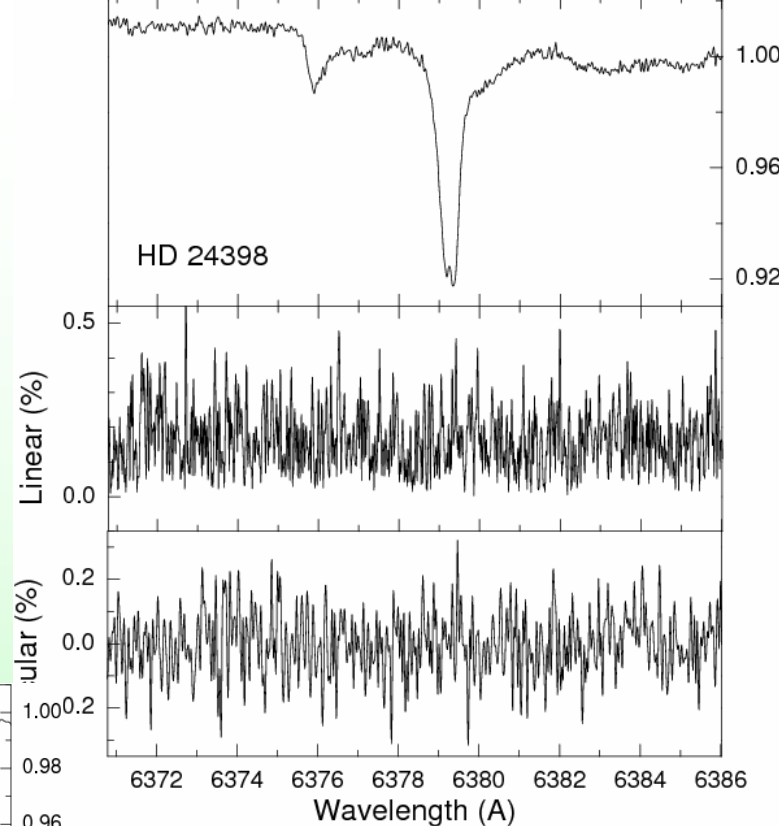
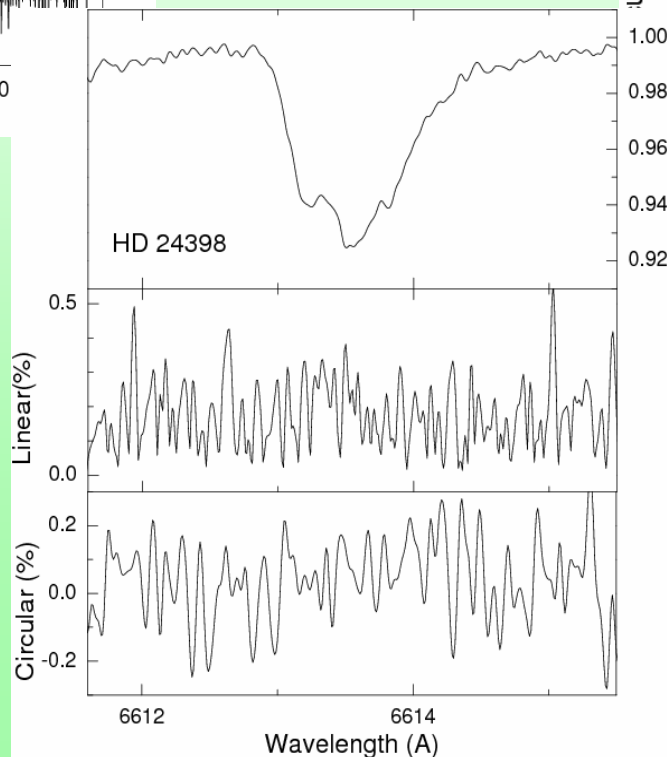
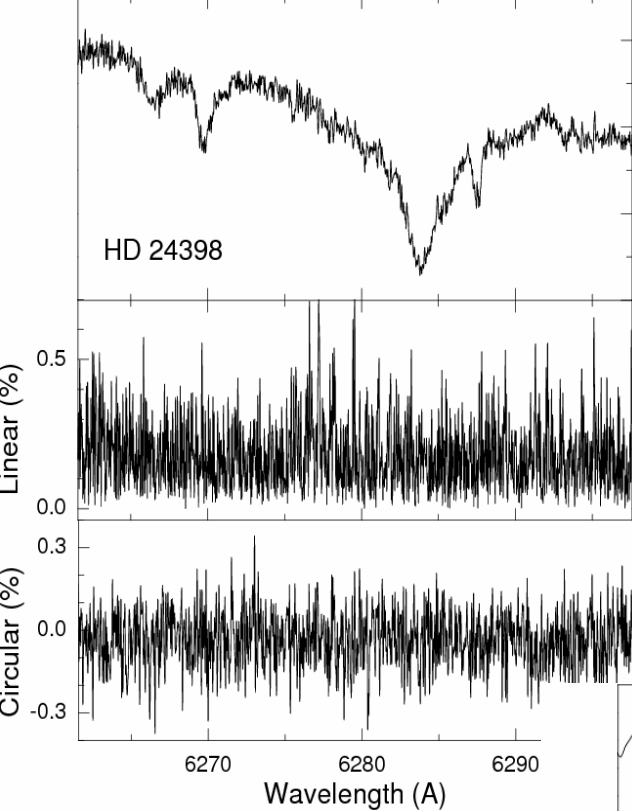


Welty & Hobbs, 2001, ApJS, 133, 345

ζ -type cloud

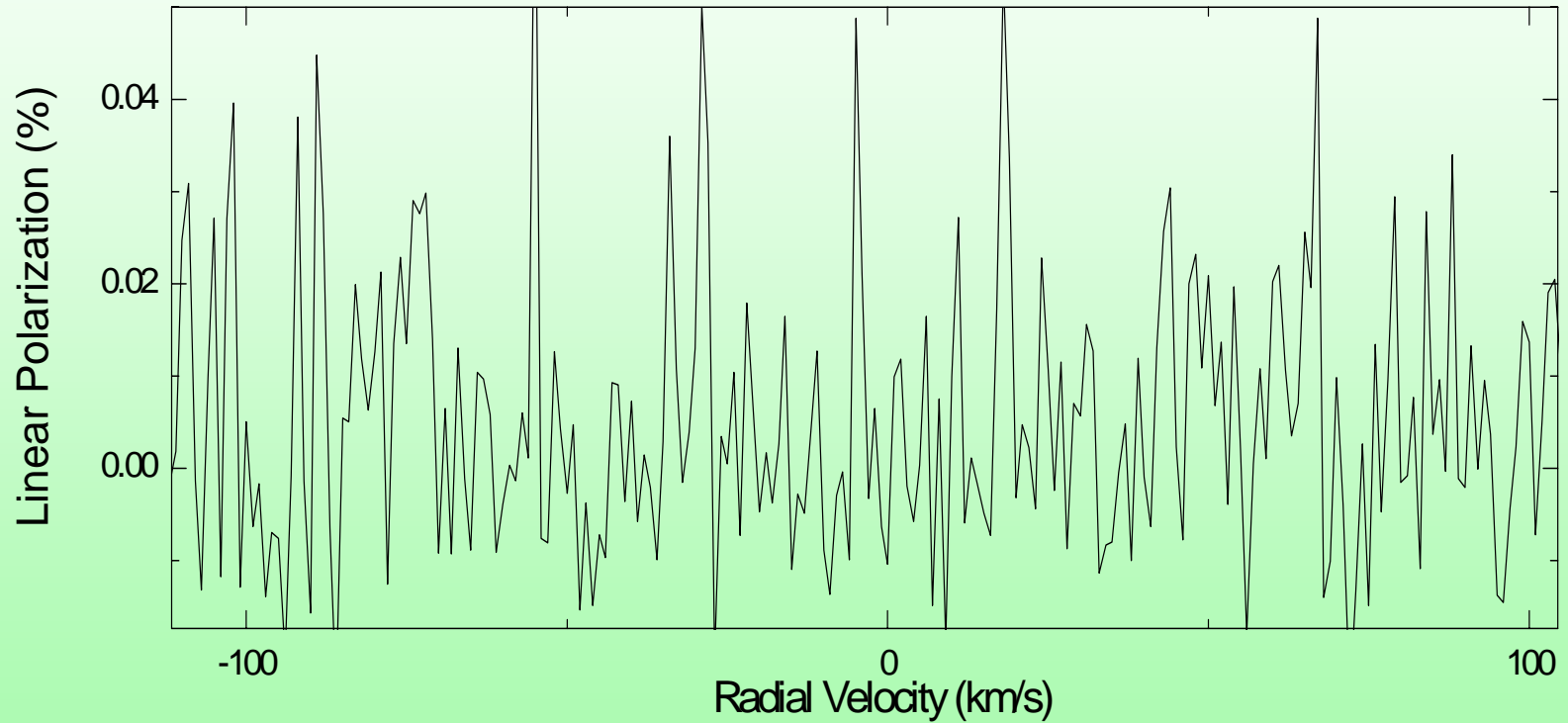






HD 24398

RMS spectrum made using 6 individual polarization spectra



THANK YOU!